

QU Carinae: Type Ia supernova in the making?

S. Kafka,¹★ R. K. Honeycutt² and R. Williams³

¹*Department of Terrestrial Magnetism, Carnegie Institution of Washington, 5241 Broad Branch Road NW, Washington, DC 20015, USA*

²*Astronomy Department, Indiana University, Swain Hall West, Bloomington, IN 47405, USA*

³*Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA*

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ABSTRACT

Variable Na I absorption lines have been reported in a number of Type Ia supernovae (SNeIa). The presence of this circumstellar material suggests that cataclysmic variables (CVs) with a giant donor star may be the progenitors of these SNeIa. We present echelle spectra of the CV QU Carinae which strengthen the connection between CVs of the V Sge class, the accretion wind evolution scenario, variable wind features, variable Na I absorption and SNeIa. This thread not only provides insight into the spectral peculiarities of QU Car, but also links SNeIa as a class with their parent systems.

Key words: novae, cataclysmic variables – supernovae: general.

1 INTRODUCTION

Although the progenitors of Type Ia supernovae (SNeIa) are still controversial, it is accepted that at least some of the events originate from detached or semidetached binary star systems in which at least one of the two components is a massive C–O white dwarf (WD). In semidetached systems, either via Roche lobe overflow of the companion or via a wind, the WD accumulates H or He-rich material which is then burned to C and O. Under the right conditions (which are primarily controlled by the stability of the mass accretion rate on to the WD), the net WD mass reaches the Chandrasekhar mass limit initiating a series of thermonuclear reactions eventually leading to a SNIa. Although outlining this scenario seems to be rather straightforward, its specifics are far from being understood or defined, especially the nature of the mass-losing star.

Among the most promising SNIa progenitors are the V Sagittae (V Sge) type cataclysmic variables (CVs; Steiner & Diaz 1998). These are semidetached binaries consisting of a WD and a giant or main-sequence donor star, with high mass transfer rates ($\sim 10^{-7}$ to $10^{-5} M_{\odot} \text{ yr}^{-1}$). This rate allows for stable nuclear burning on the surface of the WD, which can take place when a relatively massive ($0.7\text{--}1.2 M_{\odot}$; van den Heuvel et al. 1992) WD accretes near Eddington rates ($\sim 10^{-6} M_{\odot} \text{ yr}^{-1}$).

The process of a WD accumulating mass in a V Sge-type system, gradually reaching the critical Chandrasekhar mass, is described in detail by the accretion wind evolution (AWE) scenario, coined by Hachisu & Kato (2003a, and references therein). AWE successfully reproduces the long-term light curve of the prototype of the category V Sge and its Large Magellanic Cloud twin RX J0513.9–6951 (Hachisu & Kato 2003b) accounting for the bright and faint states of the systems and the transitions between them. When the accumulated envelope on the WD reaches a critical

mass, the WD atmosphere expands, generating a massive wind with $\dot{M}_{\text{wind}} \sim 10^{-7} M_{\odot} \text{ yr}^{-1}$. This wind drives a Kelvin–Helmholtz instability at its interface with the disc, peeling off the surface layer of the disc, which completely obscures soft X-rays from escaping from the binary. With time (and accumulated material) \dot{M}_{wind} increases, reaching $\sim 10^{-5} M_{\odot} \text{ yr}^{-1}$, in which case the mass of the WD envelope is further reduced by wind mass losses. The chromosphere of the companion is then eroded by the wind to the point where the donor star shrinks inside its Roche lobe and accretion (through L1) temporarily stops. The wind mass loss is then decreased, eventually reaching zero, and the system transitions to a faint state. During this faint state, soft X-rays emerge. The system gradually recovers to the bright state when the donor star regains contact with its Roche lobe and accretion commences; the cycle starts again.

Other than the prototype, there are four suspected members in this category (Steiner & Diaz 1998), but other than V Sge itself, their properties are not well known. Recently, the Galactic CV QU Carinae (hereafter QU Car) was suggested to also be a V Sge star (Kafka, Anderson & Honeycutt 2008, hereafter Paper I). The system displays strong and pronounced outflows in the ultraviolet (UV) and in optical carbon lines, reaching velocities of at least 5700 km s^{-1} (Drew et al. 2003; Paper I). At the same time, strong [O III] 5007 emission confirms the presence of a circumbinary nebula (Paper I). QU Car, being the brightest member in the V Sge category, may be a galactic twin of supersoft X-ray sources. Here, we present new high-resolution spectra of QU Car, revealing the highly variable nature of the system on time-scales of years, and we discuss its SNIa progenitor status.

2 OBSERVATIONS

Spectra were obtained with the Echelle Spectrograph on the du Pont 2.5-m telescope of the Las Campanas Observatory¹ during 6 nights

★E-mail: skafka@dtm.ciw.edu

¹ <http://www.lco.cl/>

Table 1. Log of observations.

Date of obs.	No observation	Exp. time (s)	Notes
E1: 2010-08-19	2	1800	Clear
E2: 2010-08-20	3	1800	Clear
E3: 2010-08-21	2	2000	Clear
E4: 2010-08-24	3	2000/2400	Partly cloudy
E5: 2011-02-16	2	1500	Clear
E6: 2011-02-19	2	1500	Clear

of observations. The Echelle Spectrograph provides wavelength coverage from 3700 to 9000 Å at a typical resolution of $\sim 30\,000$ (using the 1-arcsec slit). Spectra of a ThAr lamp were obtained for wavelength calibration at the position of the telescope, just before and just after each stellar exposure. A short log of observations is presented in Table 1. We use the designations E1–E6 for each of the six observing epochs and will refer to them as such for the remainder of this work. With this set-up, we obtained a total of 14 spectra of QU Car. For data processing and reductions we used IRAF’s² echelle package. Furthermore, radial velocity (RV) standard stars were observed each night and were used to confirm that our velocities for spectral features in QU Car had no systematic errors. No spectrophotometric or telluric standards were obtained during our observing runs; therefore, telluric features are present in the reduced spectra. All velocities discussed in this paper are heliocentric.

The weather was clear for most of our run. Passing clouds compromised the signal-to-noise ratio of the E4 spectra despite our increased exposure time; therefore, we will not use the E4 spectra in our analysis.

3 DISCUSSION

In V Sge, the optical photometric bright state lasts for about 170 d and the faint state for about 130 d (Robertson, Honeycutt & Pier 1997; Šimon & Mattei 1999). Fig. 1 presents part of the long-term light curve of QU Car: visual data are from the data base of the American Association of Variable Star Observers (AAVSO)³ and the V magnitude data are from the All Sky Automated Survey (ASAS; Pojmanski 1997).⁴ The full (more than 20 years) AAVSO visual light curve is also presented and discussed in Paper I. The ASAS V magnitude light curve has a detection limit of $V \sim 14$ mag. Sporadic coverage does not allow for a detailed study of the short-term behaviour of the system; however, the light curves do indicate the occasional presence of ‘faint’ states, where the brightness drops below visual 12 mag (also see Paper I). Albeit short, those faint states can last for ~ 100 d before the system recovers to its usual V ‘bright’ magnitude of ~ 11.5 . The inset in Fig. 1 is a close-up of the light curves at the times of our 2006/2007 observations, also presented in Paper I, where the system was in its bright $V \sim 11.5$ mag, exhibiting low-amplitude erratic variations which are normal for nova-like CVs. No light curves exist for the times of the echelle data of 2010–11. Therefore, we will use the spectral appearance of QU Car from Paper I as a guide on the behaviour of the system in its bright state

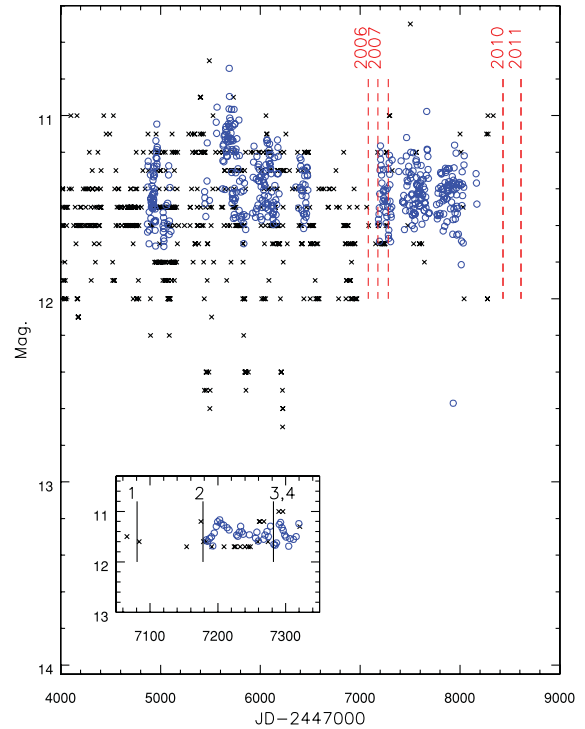


Figure 1. Visual AAVSO and V magnitude ASAS light curves of QU Car. This plot demonstrates that bright and faint states are present in the long-term optical behaviour of QU Car, although the transitions are not well defined due to lack of continuous monitoring. The four epochs of the 2006–07 low-resolution spectroscopic observations of Paper I are highlighted in the inset to demonstrate that QU Car was at its ‘bright’ magnitude at the time of the relevant observations. Photometric data were not available during our 2010–11 observations.

and as a comparison for the new spectroscopic data. Overall, the characteristics of the bright states of QU Car have been explored spectroscopically in Gilliland & Phillips (1982), Drew et al. (2003) and Paper I.

As discussed in Paper I, in the bright state the system exhibits variable and irregular Balmer emission lines, He I and He II emission and one of the strongest Bowen blend (C III/IV/O II at 4640–4650 Å) among all CVs. The broad C IV 5807 Å emission is accompanied by pronounced P Cygni profiles reaching velocities of ~ 5700 km s^{−1}. Excess of CNO-processed material in the optical spectra likely originates from a wind expelled from the binary’s C–O WD, in agreement with the AWE scenario and with relevant UV observations (Drew et al. 2003). It is critical to stress that such a wind, carrying away the outer layers of the disc and of the donor star, is necessary in the AWE scenario in order for the WD to eventually reach the Chandrasekhar mass; otherwise, the WD envelope expands and forms a circumbinary envelope, eventually leading to merging of the two stellar components instead of a SNIa (Hachisu, Kato & Nomoto 2008). The presence of [O III] 5007 Å and [N II] 6584 Å at all four epochs of the bright state in Paper I is the imprint of a nebula in the vicinity of the binary and a relic of material expelled from it.

In following sections we will discuss the time variability of the QU Car spectrum in our high-resolution data.

Gilliland & Phillips (1982) provided an orbital period of 10.9 h for QU Car; however, this period was not confirmed in Paper I, and there is not yet a reliable ephemeris for the binary. The spectral variability discussed in this paper may be due to the varying viewing angle of the different line-forming regions as the system revolves. On the

² IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

³ Observations from the AAVSO International Database (Henden A. A., private communication): <http://www.aavso.org>.

⁴ <http://www.astrouw.edu.pl/asas>

other hand, the changes may be due to variations in the strength and motion of the wind. We argue for the latter interpretation, but we admit that the lack of an ephemeris precludes our testing for orbital dependences.

Since we do not have flux-calibrated echelle spectra in hand, we use IRAF's *continuum* function to normalize the spectra in each echelle order. This allows comparison of the spectra at different epochs; however, it smeared out features that can be as wide as the echelle orders themselves. In following subsections we will discuss various emission and absorption features separately.

Finally, to address the origin of various components, we need to assess the accuracy with which we can measure the radial velocities of the lines and the possible systematic errors affecting their respective values. It is always possible that instrumental effects (echelle spectrograph flexure) introduce pixel shifts and pseudo-variations in the radial velocities of the lines. Taking comparison lamp spectra before and after each object exposure is a common procedure to minimize this effect; nevertheless, an independent confirmation is necessary. In the spectrum of QU Car, there are two diffuse interstellar bands (DIBs) at 5870 and 5797 Å at adjacent blue orders, the oxygen *A* and *B* bands and the telluric bands at red orders. We have used the narrow DIB features to ascertain the accuracy of our wavelength calibrations inasmuch as the RV shifts in the Na I D absorption occur at different temporal epochs and could conceivably be due to slight changes in the wavelength calibration due to the instrument configuration. Measurement of the DIBs radial velocities confirms the accuracy of our RV determinations within an rms error of $\sim \pm 0.10 \text{ km s}^{-1}$.

3.1 Emission lines

Table 2 lists the equivalent widths (EWs) of the most prominent features in our echelle spectra. By comparing to the emission line strengths in 2006/2007 (table 3 of Paper I) we see that the new EWs are in general consistent with those measured in 2006/2007 using lower resolution data. Both the 2006/2007 EWs and the new data are characterized by occasional changes of over two times between nights, and sometimes within a night, mixed with intervals of relative stability. Note in particular that in 2006/2007 the Bowen blend was nearly constant over 6 months at $\text{EW} \sim 3.2 \text{ Å}$, which is a little weaker than in 1979/1980 (Gilliland & Phillips 1982). In our 2010 spectra this feature is much weaker and has practically disappeared by E3. The Balmer lines are also weak at this time, but He II is stronger than the 2006–11 average. Apparently these

various emission lines originate in quite different regions of the system, which can vary independently.

Figs 2 and 3 show the line profiles of some of the emission lines, where the higher spectral resolution of the echelle data reveals considerable details compared to earlier studies. At times the H α profile resembles the line profiles in V Sge (Robertson et al. 1997; Gies, Shafer & Wiggs 1998), with multiple components whose relative strengths vary. Note that the H α profile sometimes shows weak blueshifted absorption at a velocity of $\sim -300 \text{ km s}^{-1}$. This wind-induced P Cygni feature was also seen in the H α profiles of Paper I.

For all four of the 2011 spectra (E5 and E6) the H α emission line is accompanied by seven narrow weak absorption features. These

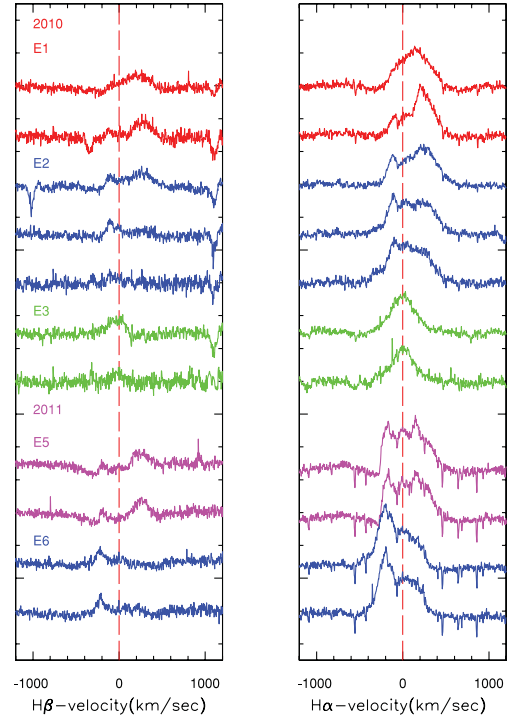


Figure 2. Balmer lines for epochs E1–E6 (E4 is missing, since the relevant data are not used in this work). The Balmer line profiles seem to be smoother in 2010, with multiple absorption components in 2011. A stationary redshifted emission component at $\sim -135 \text{ km s}^{-1}$ is representing stationary circumbinary gas. Finally, the blueshifted and redshifted absorption troughs in the 2010 H β line likely represent material expelled from the binary.

Table 2. EWs^a (in Å) of the emission features of QU Car.

Date obs.	Epoch	HJD – 245 5428	Bowen C III/N III/O II	He II 4686 Å	H β	He I 5876 Å	H α	He I 6678 Å	He I 7065 Å
2010	E1	0.501	–2.23	–1.47	–0.77	–0.18	–2.72	–1.12	–0.38
		0.556	–1.78	–1.98	–1.25	–0.85	–3.11	–1.38	–0.36
		1.399	–0.70	–1.89	–0.79	–0.10	–3.23	–0.73	–1.13
		1.546	–1.00	–2.13	–0.61	–0.39	–3.78	–0.12	–1.20
	E3	1.526	–1.03	–5.14	–	–0.39	–2.98	–0.58	–1.62
2011	E5	2.479	–0.06	–2.75	–0.63	–0.99	–2.26	–	–
		2.503	–1.00	–3.05	–0.28	–	–2.28	–	–
		180.757	–3.39	–1.03	–0.57	–0.44	–4.07	–0.66	–
		180.775	–3.31	–0.74	–0.69	–0.24	–3.01	–0.80	–
	E6	183.831	–3.42	–1.08	–1.43	–0.26	–3.75	–1.42	–0.98
		183.849	–2.34	–1.35	–1.01	–0.38	–4.47	–1.18	–1.16

^aEW measured error: 0.02 Å.

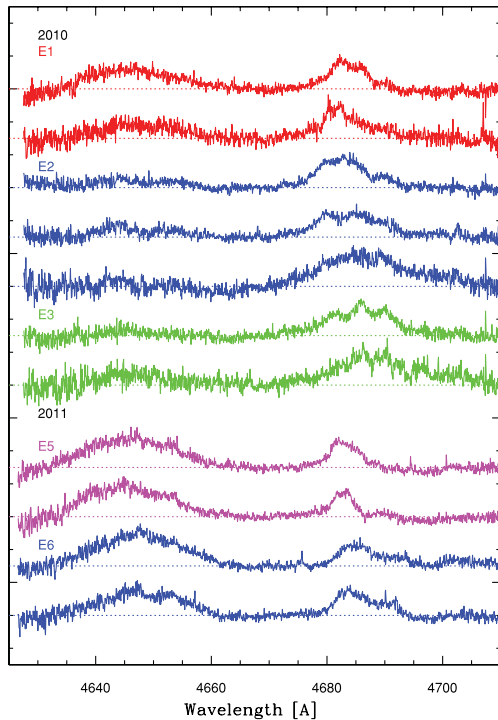


Figure 3. Same as Fig. 2, but for the Bowen blend and He II 4686 emission. The dashed horizontal lines guide the eye to the normalized continuum. Note the attenuated Bowen blend lines in 2010 with respect to their appearance in 2011.

features are identical in these four spectra (separated by 3 nights) and do not appear in any of the 2010 spectra (E1–E3). They do not agree with the wavelengths of telluric water features (e.g. Vince & Vince 2010), nor do they appear in our standard stars (or any of the stars observed the same night in neighbouring parts of the sky, bracketing the QU Car observations). In the absence of an accurate ephemeris, we have few clues to the physical locations of those lines.

The fact that the EWs of He II and He I do not follow changes in the Balmer and Bowen blend suggests that those lines have different origins in the system (as expected from their different excitation potentials). Using RV arguments, Gilliland & Phillips (1982) argue that C III 4650 Å should be the dominant species in the Bowen blend. Furthermore, UV spectra (Drew et al. 2003) indicate that carbon is overabundant in QU Car. A carbon overabundance in QU Car is also supported by the presence of the strong C IV 5801, 5812 Å emission, in the Gilliland & Phillips (1982), Drew et al. (2003) and Paper I optical spectra. Drew et al. (2003) argue that this carbon enhancement should originate from the envelope of the donor star, suggesting that the star is an early-R type. However, the observed variations in the Bowen blend and in C IV in our spectra argue against such an interpretation. Emission from the secondary star could not have been diminished in one epoch of our observations but still be present at a different one. In 2010 both the Bowen blend and C IV are significantly reduced in strength; this is an indication that the lines are variable in nature. A plausible interpretation is that carbon originates from the atmosphere of the WD of the binary, which is attenuated in the faint state. In 2011, when the C III/N III/O II line strength of the binary returned to its ‘bright’ state value, weak C IV 5801, 5812 Å emission is also present.

In our 2006/2007 spectra, we also detected variable [O III] 5007 Å and [N II] 6584 Å emission, indicative of the presence of a nebula

around the system. The [O III] 5007 Å lines appeared to have two components, ranging in velocities from -500 to 370 km s $^{-1}$ – probably representing the front and back of an expanding shell or an outflow. In our new echelle data, only one (weak) component of the [O III] 5007 Å line is present; its strength is reduced with time (perhaps with orbital phase) between E1 and E2. In 2011, the [O III] 5007 Å emission is present only during E5, with no accompanying [N II] 6584 Å emission. This erratic variation in the oxygen emission line strength does not seem to be associated with variations in the strength of the Bowen blend. However, we are in need of a better ephemeris for the system in order to correlate the observed oxygen emission to a specific location and process on the binary.

3.2 H β components: evidence for circumstellar material?

We now turn our attention to the mysterious absorption components bracketing H β in 2010. Having excluded contamination in our spectra (flat-fielding artefacts, data reduction faux pas or background reduction residuals), the features are convincingly real. Since there are no strong telluric features at this part of the spectrum (nor is there any chance for this absorption to come from the interstellar medium, especially considering its variable nature and appearance), the features should originate from the binary. The red component is stationary at ~ 1110 km s $^{-1}$. The blue component is present only during two subsequent nights of our observations (second spectrum in E1 and first spectrum in E2, at -370 and -1050 km s $^{-1}$, respectively). Considering that our exposure times are 2000 s, we are surprised to not detect this red component in all spectra in E1 and E2.

Among CVs, those features are unique to QU Car. However, similar absorption lines sometimes appear in other accreting sources: complex blueshifted and redshifted absorption components emerge in the Balmer lines of Herbig Ae/Be stars (e.g. Guimarães et al. 2006). In this case, the redshifted absorption is usually interpreted as being due to material accreted on to the star from its inner disc (sometimes via magnetic fields) and the blueshifted one as being due to mass loss (outflows). Both processes can happen simultaneously (e.g. Natta, Grinin & Tambovtseva 2000). Taking the Ae/Be work as a starting point, we could attribute the blue components to ejecta from QU Car, while the redshifted components originate on an inflow from the circumstellar medium. It is difficult to imagine a mechanism that allows for both processes simultaneously, unless the flows are magnetically controlled (as is the case in many T Tauri stars). However, there is no evidence that QU Car contains a highly magnetic WD (i.e. no Zeeman or cyclotron features). Furthermore, it seems unlikely that this material represents a jet. Jet-like features in CVs sometimes appear in dwarf novae in outburst or in some novae, both times representing material ejected from the binary during explosions (e.g. Cowley et al. 1998). In semidetached binaries, optical jets also appear in symbiotic stars and in low-mass X-ray binaries as *emission* components moving away from Balmer lines with velocities reaching 4000 km s $^{-1}$ (e.g. Cowley et al. 1998). Perhaps the gas motions in QU Car are so turbulent and complex that portions of the flow can be both redshifted and blueshifted in front of the same continuum source. It is unfortunate that we lack optical photometry (or any light curves) at the time of our spectroscopic observations (stressing again the importance of having simultaneous multiwavelength observations for a comprehensive understanding of QU Car).

3.3 Na I D absorption: connection with the Type Ia supernovae?

The unusual appearance of the Na I D lines is revealed in Fig. 4: in 2010, each component consists of weak blueshifted emission accompanied by redshifted absorption! Those lines (and the adjacent

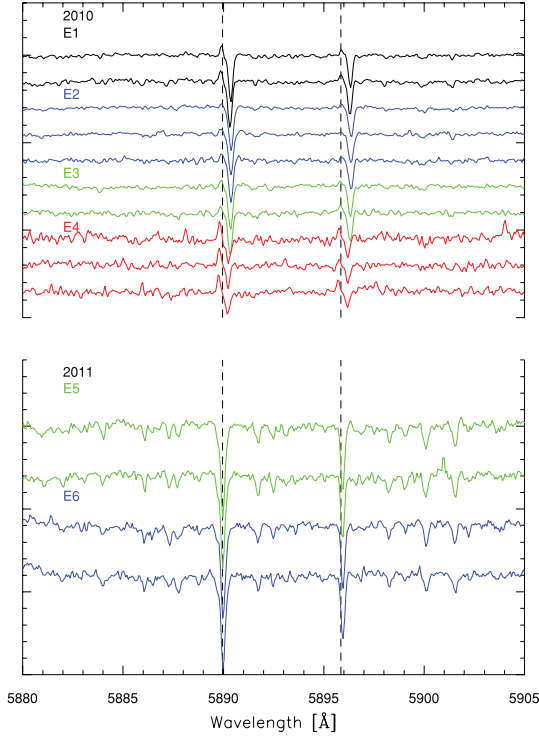


Figure 4. Same as Fig. 3, but for the Na D line region.

He I 5876 emission) are at the centre of the order during all our observing runs, ruling out RV variations introduced by pixel shifts due to poor echelle order distortion correction at the edges of each order. The EW and RV values are given in Table 3, where missing entries indicate that the line was not detected. The emission components are very narrow and appear slightly blueshifted with respect to the line's rest wavelength. Furthermore, this emission is quite variable and present only during 2010. No correlation of the line strength and RV was found. In CVs, the Na I D lines are usually in absorption, arising from the photosphere of the donor star. Sodium is usually ionized in accretion discs; its presence in emission in QU Car is indicative of lower than usual temperatures. This suggests a different optical state for the system, perhaps similar to V Sge's 'faint state', where the mass transfer rate is significantly lower, in agreement with the reduced strength of the Balmer and the Bowen blend line profiles. We are intrigued by the 'stationary' (constant velocity) nature of this Na I D emission component over the 4 nights of our observations in 2010. If this line has an origin on the accretion disc or on the hotspot, it may indicate that the QU Car's disc outer edge reaches the centre of mass of the binary, in agreement with QU Car being an unusually bright object.

The RVs of both Na D absorption components are constant at $\sim -7 \text{ km s}^{-1}$ during E1–E3. 8 months later these RVs are again constant but at $\sim -13 \text{ km s}^{-1}$ during E5–E6. Both components of the Na I D are well displaced from their rest wavelength as they accelerate between the 8 months of observations. In the bright state, variable Na I D is also present (e.g. Paper I); however, with the lower resolution of Paper I it was not possible to resolve individual absorption components. Furthermore, in our echelle spectra, we cannot kinematically associate any other emission line components with the absorption features.

Since neither component of the Na I D doublet is at the local standard of rest velocity, we explore the possibility that they are

Table 3. RV (in km s^{-1}) and EW (in \AA) measurements of the Na D lines.

	HJD – 245 5428	RV (km s^{-1})	EW (\AA)	RV (km s^{-1})	EW (\AA)
		Na D1 emission	Na D1 emission	Na D1 absorption	Na D1 absorption
E1	0.501	16.66	–0.02	–6.98	0.14
	0.556	18.73	–0.03	–5.01	0.15
E2	1.400	16.26	–0.01	–7.93	0.14
	1.546	–	–	–7.52	0.13
E3	1.526	–	–	–7.62	0.11
	2.479	17.42	–0.01	–6.72	0.14
E5	2.503	18.90	–0.02	–5.80	0.12
	180.757	–	–	–13.37	0.15
E6	180.775	–	–	–13.15	0.16
	183.831	–	–	–13.71	0.16
	183.849	–	–	–13.76	0.16
		Na D2 emission	Na D2 emission	Na D2 absorption	Na D2 absorption
E1	0.501	16.30	–0.02	–6.75	0.09
	0.556	17.75	–0.02	–5.50	0.09
E2	1.400	15.80	–0.01	–8.52	0.10
	1.546	–	–	–7.96	0.10
E3	1.526	–	–	–8.05	0.09
	2.479	18.79	–0.02	–6.75	0.09
E5	2.503	18.84	–0.03	–6.04	0.09
	180.757	–	–	–13.22	0.09
E6	180.775	–	–	–13.25	0.08
	183.831	–	–	–13.77	0.08
	183.849	–	–	–13.40	0.09

^aRV $\pm 0.10 \text{ km s}^{-1}$; ^bEW $\pm 0.02 \text{\AA}$.

interstellar in origin, following the Galactic rotation at the location of QU Car. Drew et al. (2003) estimate the galactic rotation velocity at ~ 2 kpc along the line of sight to QU Car to be -18.3 km s^{-1} and attribute Fe II 1608 Å and Si II 1250 Å blueshifted line components to interstellar absorption. As a side note, the minimum distance to QU Car is suggested to be 500 pc (Linnell et al. 2008, and references therein), which would imply less interstellar absorption. However, for our purposes, the distance to QU Car and amount of interstellar absorption is irrelevant since we are discussing RV *variations* in the absorption lines, not their absolute values.⁵ In this case, the variable velocities of Na I D components between 2010 and 2011 argue against an interstellar origin. Moreover, Na I D absorption could originate on the donor star of the binary. Because of the absence of other photospheric identification features from this star (e.g. Fe I, Ca I, Mg II, Ti or even Balmer absorption) we rule out this possibility. In CVs, Na I absorption lines can be generated in nova explosions (e.g. Shore et al. 2011); however, QU Car has no indications of recent eruptions.

A possible scenario comes from the supernova community: Patat et al. (2007) describe variable Na I D absorption in the vicinity of SN 2006X, originating from circumbinary clouds, heated by the SN and interacting with the explosion ejecta. Patat et al. (2007) argue that the Na I was present in the SNIa progenitor as circumbinary material. UV radiation from the SNIa explosion ionized this material, which slowly recombined post-SNIa to produce the observed variable Na I D absorption. The absence of Ca II H&K absorption lines was attributed to the lack a radiation field hot enough to significantly ionize Ca.

Variable Na I D absorption has now been detected in a group of SNeIa (Simon et al. 2007, 2009; Sternberg et al. 2011), favouring the single degenerate scenario (C–O WD + giant donor star) for a SNIa progenitor. Circumstellar clouds continuously replenished by a wind from the giant donor star provide a plausible mechanism for the formation of Na I D lines in QU Car. Therefore, our detection of variable Na I D lines around QU Car provides, for the first time, a bridge between this class of SNe and their elusive progenitors.

4 FINAL REMARKS

Williams et al. (2008) has reported variable transient heavy element absorption lines after a nova explosion, likely originating from interactions between circumbinary material and the nova ejecta. Density enhancements of circumbinary material have also been detected around the recurrent nova RS Oph (Iijima 2008; Patat et al. 2011), likely originating from expanding material from the donor star, urging observers to conduct similar studies. QU Car is the first system in which such material is discovered and characterized, providing evidence that high \dot{M} CVs are the likely progenitors of SNeIa. We have presented two epochs of high-resolution observations of this system, attributing emission from the carbon-dominated Bowen blend to the WD of the binary. Strong and variable absorption components to the H β emission lines attest to the presence of circumbinary material that is receding from the binary. Furthermore, the presence of blueshifted Na I D absorption components is interpreted as being due to circumbinary material likely removed from the donor star by a strong wind. Such material has also been detected around SNeIa and is now a signature of the donor star of some SNeIa. Most importantly, this material corroborates the valid-

ity of the AWE scenario for mass increase on to the WD, eventually leading to a SNIa explosion.

The variability of the spectrum of QU Car presents many mysteries and challenges, ranging from the unidentified absorption features near the Balmer lines to the multiple components of the strong emission lines to the redshifted Na I D absorption line profiles with blue emission components. To add to the confusion, most of these effects are highly variable. It is frustrating that our attempts to understand these phenomena have not been broadly successful, but we do wish to draw attention to these features in order to encourage further work. Also, we wish to emphasize that the lack of a reliable orbital ephemeris greatly complicates the interpretation of widely spaced spectral features such as the ones presented here. Our present study documents for the first time a phase where accretion is diminished in this system in 2010. More importantly, this new study ties together the chasm between observations of SNeIa and their likely progenitors, providing a viable candidate of a SN in the making. QU Car is indeed a unique system, even among famously divergent CVs.

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⁵ It is also possible that an interstellar component is present and unresolved in the observed lines; however, it will simply introduce an error in the estimate of the line EW.